

Astron. Astrophys. Suppl. Ser. **46**, 185-191 (1981)

Photometry and elements of GW Geminorum

P. Broglia and P. Conconi

Osservatorio Astronomico, via E. Bianchi 46, 22055 Merate, Italy

Received December 30, 1980, accepted June 6, 1981

Summary. — Photoelectric observations of the eclipsing system GW Gem made during the years 1978 and 1979 enable us to derive nine epochs of minimum light. The period of the binary star appears to be constant. The light curves are analysed referring to the Roche model. The brighter component appears to be an A8 star and the secondary G3-G4 subgiant fills its Roche lobe.

Key words : Eclipsing binary — photometric solutions.

1. Light curves and period. — GW Gem = BD + 27°1494 was discovered to be an Algol-type variable by Hoffmeister (1949). On the basis of photographic material obtained at Sonneberg and Babelsberg Observatories which began in 1928, Löchel (1955) derived several epochs of minimum. Some other photographic epochs have been derived by Kippenhahn (1953) and by Romano and Perissinotto (1968). On the basis of all this material, 48 minima altogether, Romano and Perissinotto calculated the light elements :

$$\text{Min I} = \text{helioc. JD } 2425645.572 + 0.65944354 n .$$

During the following years GW Gem has been monitored assiduously by the Observers of BBSAG (1973-1978) which on the basis of visual measurements calculated 27 epochs of minimum using the tracing paper method. However, as far as we know, no photoelectric light curve has been obtained so far for this relatively bright and short period system.

The observations studied in this article were made through *B* and *V* filters, a Lallemand photomultiplier and a Gardiner type integrator at the 102 cm reflector of the Merate Observatory, during the years 1978 and 1979. The integration time was set to 20 s. The variable was observed differentially relative to the star BD + 28°1494 ; the check star was BD + 27°1497. The extinction coefficients were determined on some nights ; when it was not possible, mean values were used. The extinction corrections exceptionally came out to 0^m.005 for *B* and *V* measurements, but for the most part they were quite negligible.

At first sight the brightness of the comparison relative to the check star appeared to be constant. A total of 56 *B* and 54 *V* differential magnitudes $m_{\text{comp}} - m_{\text{check}}$ gave indeed :

$$\begin{aligned} \Delta B &= - 0^{\text{m}}.285 \pm .001, \\ \Delta V &= - 0^{\text{m}}.018 \pm .001 \text{ m.e.} \end{aligned}$$

Send offprint requests to : P. Broglia.

The corresponding standard deviations for a single Δm , 0^m.010 and 0^m.009 respectively in *B* and *V* conformed to the usual observing conditions in Merate. However when computing the photometric solutions of GW Gem, the residuals of the observations obtained during some nights showed clearly a small and periodic trend. Since no second check star was used and since the number of observations on GW Gem widely exceeds the number of Δm between the comparison and check star, the observations of GW Gem itself better support the variability of BD + 28°1494. Effectively, after comparing both with GW Gem, cleared from the eclipse effects, and with BD + 27°1497, the comparison star turned out to be a δ Scuti star with an overall amplitude of about 0^m.035 at the most (Broglia and Conconi, 1979), but which can also decrease below the noise level of measurements, as we have discussed in a separate paper (Broglia and Conconi, 1981). This fact produced some complications when solving the light curves, as will be said in the next section.

The reference star was tied to the standard *UBV* system on three nights. We obtained :

$$\text{BD} + 28^{\circ}1494 : \quad V = 9^{\text{m}}.32 \quad B-V = + 0^{\text{m}}.21 .$$

The *V* value refers to the mean brightness of BD + 28°1494. Taking into account the δ Scuti variation, its incertitude is evaluated at 0^m.02 at the most, whilst the colour is quite correct.

During fourteen nights we obtained altogether for GW Gem 1252 differential magnitudes in *B* and 1272 in *V*, listed in tables I and II. By fitting a second degree parabola to the lower part of each minimum we derived five epochs of primary and four of secondary eclipse, listed in table III. The quoted errors relate to the mean of *B* and *V* values. Taking into account only the primary minima we calculated the following ephemeris :

$$\text{Min I} = \text{helioc. JD } 2425645.5938 + 0.659443504 n \quad (1)$$

$$\begin{array}{ccc} \pm & 5 & 17 \text{ m.e.} \end{array}$$

To detect eventual changes in the period we have considered moreover the photographic and visual timings of light given in the literature from which we calculated the normal epochs given in table III with the corresponding mean errors, to have their precision valued. After weighting conveniently, all the instants of minimum lead to the following ephemeris :

$$\text{Min I} = \text{helioc. JD } 2425645.5798 + 0.659444013 n \\ \pm 40 \quad \quad \quad 146 \text{ m.e.}$$

The residuals O-C are given in table III. The period including all observations over 50 years would be slightly longer than the value determined from photoelectric data alone. Adding a quadratic term to the ephemeris, a reduction of the mean residual of only four percent is obtained, so on the basis of the meagre material available the period of GW Gem can be considered constant.

The observations were put in phase by means of the ephemeris (1). From the light curves given in figure 1 we deduce :

	Max	Min I	Min II
V	10 ^m 485	11 ^m 450	10 ^m 685
B-V	+ 0.27	0.33	0.24

2. Photometric solutions and concluding remarks. —

A preliminary analysis of the light curves according to the Russell model was performed by means of the Irwin (1947) differential correction method in order :

1) to separate the eclipse changes from the δ Scuti variation ;

2) to evaluate the influence on the solutions of the variability of the comparison.

Harmonic analysis of the residual derived in the photometric solutions were performed, separately for each night and colour. The periods and amplitudes, non constant, of the δ Scuti variation are given elsewhere (Broglia and Conconi, 1981). Using these values the single measurements of GW Gem have been corrected afterwards for the Scuti variation of the comparison, limited to the nights when the variation was detected and the clearing out for the δ Scuti effect proved to be significant in comparison with natural dispersion of the observations. About two thirds of the measurements have been corrected and the corresponding reduction factor of the dispersion was 0.75.

The influence on the photometric solutions of the δ Scuti variation appeared to be negligible because its amplitude is small compared with the eclipse changes and in most parts of the mean light curves the fluctuations cancel out. In fact correcting the out of eclipse measurements the standard deviation for a single observation with respect to the Fourier representation decreased from 0^m009 to 0^m007 in B and from 0^m008 to 0^m006 in V. A similar improvement was attained for the measurements during the eclipse. We note moreover that when solving according to the Russell model the $\sin n\theta$ terms resulted quite negligible.

Starting from these preliminary elements, solutions were then calculated by means of the Wilson and Devinney (1971) programmes, updated to 1978. In order to reduce the computing time when operating with the differential corrections programme we make the

following change. According to the Wilson and Devinney notations the flux sorting in the observing direction from each areola in which both stars are subdivided is equal to :

$$\Delta F_{\lambda} = \cos \gamma \text{ DIGR}^2 \text{ sen } \theta \Delta \theta \Delta \phi / \cos \beta .$$

For a given orbital phase the factor $\cos \gamma$, where γ is the angle between the local surface normal and the line of sight, depends only on the « geometrical » parameters (i , q , Ω_1 , Ω_2); the limb darkening D depends on these parameters and on the coefficients x_1 and x_2 ; the product $\text{IGR}^2 \text{ sen } \theta / \cos \beta$ depends not only on the geometrical parameters but also on the « photometric » ones (g_1 , g_2 , T_1 , T_2 , L_1 , L_2 , A_1 , A_2 , x_1 , x_2), but not on the phase; $\Delta \theta$ and $\Delta \phi$ depend on the grid's numbers. In the original programme the quantity $\text{IGR}^2 \text{ sen } \theta / \cos \beta$ is computed for the initial set of parameters and the corresponding vector is memorized. Then the luminosity l is computed for each phase by means of the subroutine « Light ». If we want to compute the differential corrections for n parameters we must repeat n times this process after having incremented successively each of the n parameters. On the whole $n + 1$ light curves are computed and from these the derivatives $\delta l / \delta p$ belonging to each of the n parameters p are calculated.

The programme can be improved by storing in a matrix the factor $\text{IGR}^2 \text{ sen } \theta / \cos \beta$ computed for the initial set of parameters and for the n sets obtained by increasing by Δp , one after another, each of the n photometric parameters to be corrected. In this way the quantity $\cos \gamma$, corresponding to a given phase and to a given areola, can be computed in the subroutine « Light » only once for all the photometric parameters and we can derive simultaneously for each observation the ΔF_{λ} pertinent to all the sets of photometric parameters (the initial and the increased ones).

In the case of grids 25×25 and 15×15 , when correcting three photometric and three geometric parameters the computing time is reduced by 30 percent. This improvement is appreciable especially when using a small computer.

Because no spectroscopic study of GW Gem is as yet available in the literature, taking into account the fact that the eclipses are deep, in conformity with the colour at the Min II a spectral type A8 was attributed to the primary component, supposed to be a main sequence star and a temperature $T_h = 7\,700$ K was assumed according to the Flower (1977) temperature scale. The following parameters were allowed to vary : i , L_1 , L_2 , A_2 , Ω_1 , Ω_2 . Of the remaining parameters which were not adjusted during the calculations, the limb darkening coefficients were taken from Grygar *et al.* (1972), and the gravity darkening exponents were set equal to one for both stars. Since the spectroscopic mass-ratio q is unknown and the solutions show a correlation between q and i , we searched for solutions, by means of the differential corrections programme, with different assumed values for q ranging from $q = 0.32$ to $q = 0.60$ at convenient steps. In the range from $q = 0.32$ to 0.45 the filling factor $F_2 = \Omega_{\text{int}} / \Omega_2$ for the secondary component is 1.0 and in the interval 0.45 to 0.60 it decreases to 0.985 only ; therefore the secondary component fills its Roche lobe.

The best solution (with the smallest sum of the squares of the residuals) is at $q = 0.43$, as appears from figure 2 where the solutions calculated with F_2 as free parameter (dots) and those derived with $F_2 = 1$ (circles) are represented. The adopted solution calculated assuming q also as free parameter and using the programme in mode 0, according to Wilson' code, is listed in table IV together with the related standard errors pertinent to the adjusted parameters. The luminosity of the bright star is normalized to make the sum of the two components unity. The temperature $T_c = 5\,270$ K, calculated according to the Planck's law for the secondary component, favours a spectral type G3-4 for the secondary subgiant. The solutions were calculated using all the B and V individual observations together.

A more realistic evaluation of the uncertainty of the solutions and in particular of the mass-ratio q can be

derived by looking at figure 2. Assuming a 5 percent worsening of the factor R (see legend to figure 2) is significant, an incertitude of ± 0.05 can be assigned to q . The corresponding errors for the elements i , Ω_1 , A_2 , L_{1B} , L_{1V} are given between brackets in table IV.

In conclusion GW Gem appears to be a semidetached system ; the small value for the r.m.s. residual σ for single B or V observations computed in the solution (table IV) and the small dispersion of the measurements displayed in figure 1 confirms that the light curves have been stable and devoid of sensible perturbations over the two observing seasons and that the computed elements are reliable.

Acknowledgements. — The authors thank Dr. R. E. Wilson for providing the computer codes.

References

- BBSAG Observers : *BBSAG Bull.* n° 8, n° 9, n° 11 (1973) ; n° 13, n° 14 (1974) ; n° 19, n° 20, n° 21, n° 22 (1975) ; n° 26, n° 27 (1976) ; n° 31, n° 33 (1977) ; n° 36 (1978).
 BROGLIA, P., CONCONI, P. : 1979, *Inf. Bull. Var. Stars* n° 1595.
 BROGLIA, P., CONCONI, P. : 1981, *Astron. Astrophys.* **100**, 201.
 FLOWER, P. J. : 1977, *Astron. Astrophys.* **54**, 31.
 GRYGAR, J., COOPER, M. L., JURKEVICH, I. : *Bull. Astron. Inst. Czech.* **23**, 147.
 HOFFMEISTER, C. : 1949, *Erg. Astron. Nachr.* **12**, 1.
 IRWIN, J. B. : 1947, *Astrophys. J.* **106**, 391.
 KIPPENHAHN, R. : 1953, *Astron. Nachr.* **281**, 156.
 LÖCHEL, K. : 1955, *Astron. Nachr.* **281**, 156.
 ROMANO, G., PERISSINOTTO, M. : 1968, *Mem. Soc. Astron. Ital.* **39**, 430.
 WILSON, R. E., DEVINNEY, E. J. : 1971, *Astrophys. J.* **166**, 605.

TABLE I. — B observations of GW Gem.

Table with 12 columns: HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B. Contains multiple rows of astronomical data.

Table with 12 columns: HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B. Contains multiple rows of astronomical data.

Table with 12 columns: HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B. Contains multiple rows of astronomical data.

Table with 12 columns: HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B. Contains multiple rows of astronomical data.

Table with 12 columns: HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B, HEL. J. D. 2443., B. Contains multiple rows of astronomical data.

TABLE II. — V observations of GW Gem.

Main table containing photometric observations for GW Gem. Columns include Heliocentric Julian Date (HEL. J.D.), magnitude (V), and observation number. The table is organized into multiple columns for different observation periods.

Continuation of the photometric observations table, showing further data points for GW Gem across various observation dates and magnitudes.

Final section of the photometric observations table, containing the last set of data points for GW Gem.

TABLE III. — *Times of minima of GW Gem.*

n	helioc. JD	m.e.	O-C	Observer
0	2425645.569	± 0.007	-0.011	ph (6) L
1131	26391.416	.007	+ .005	ph (13) K
4325	28497.671	.006	- .004	ph (8) L
4514	28622.324	.006	+ .014	ph (7) K
7741	30750.346	.007	+ .010	ph (5) L
16045	36226.348	.011	- .011	ph (5) R,P
20555	39200.426	.003	- .026	ph (4) R,P
24440	41762.390	.002	- .002	v (6) BBSAG
25412	42403.367	.001	- .004	v (8) "
26008	42796.401	.002	+ .001	v (8) "
26474	43103.701	.004	.000	v (5) "
27097	43514.53438	.00003	+ .00011	pe B,C
27141	43543.55006	.00004	+ .00026	pe "
27142.5	43544.5411	.0005	+ .0021	pe "
27192.5	43577.5120	.0004	+ .0008	pe "
27210.5	43589.3798	.0010	- .0014	pe "
27605	43849.53175	.00006	- .00007	pe "
27646	43876.56881	.00006	- .00022	pe "
27690	43905.58451	.00001	- .00006	pe "
27721.5	43926.3566	.0003	- .0004	pe "

L = Löchel (1955) ; R, P = Romano and Perissinotto (1968).
 K = Kippenhahn (1953) ; BBSAG = BBSAG Observers (1973-1978).
 B, C = Broglia and Conconi, present work.

TABLE IV. — *Photometric elements of GW Gem.*

$i = 81^\circ.41 \pm 0.08$ p.e. (1.4)	$r_{1,pt} = 0.380$	$r_{2,pt} = 0.415$	$g_1 = 1.00$
$q = 0.436 \pm 0.004$ (.05)	$r_{1,s} = 0.361$	$r_{2,s} = 0.302$	$g_2 = 1.00$
$\Omega_1 = 3.271 \pm 0.003$ (.017)	$r_{1,b} = 0.372$	$r_{2,b} = 0.334$	$A_1 = 1.00$
$\Omega_2 = 2.751$ (contact)	$r_{1,p1} = 0.350$	$r_{2,p1} = 0.289$	$A_2 = 0.50 \pm 0.01$ (.03)
λ	x_1	x_2	L_1
0.430 μ	0.78	0.83	0.9283 ± 0.0006 (.007)
0.550	0.60	0.70	0.8936 ± 0.0008 (.008)
			$\sigma = 0.008$

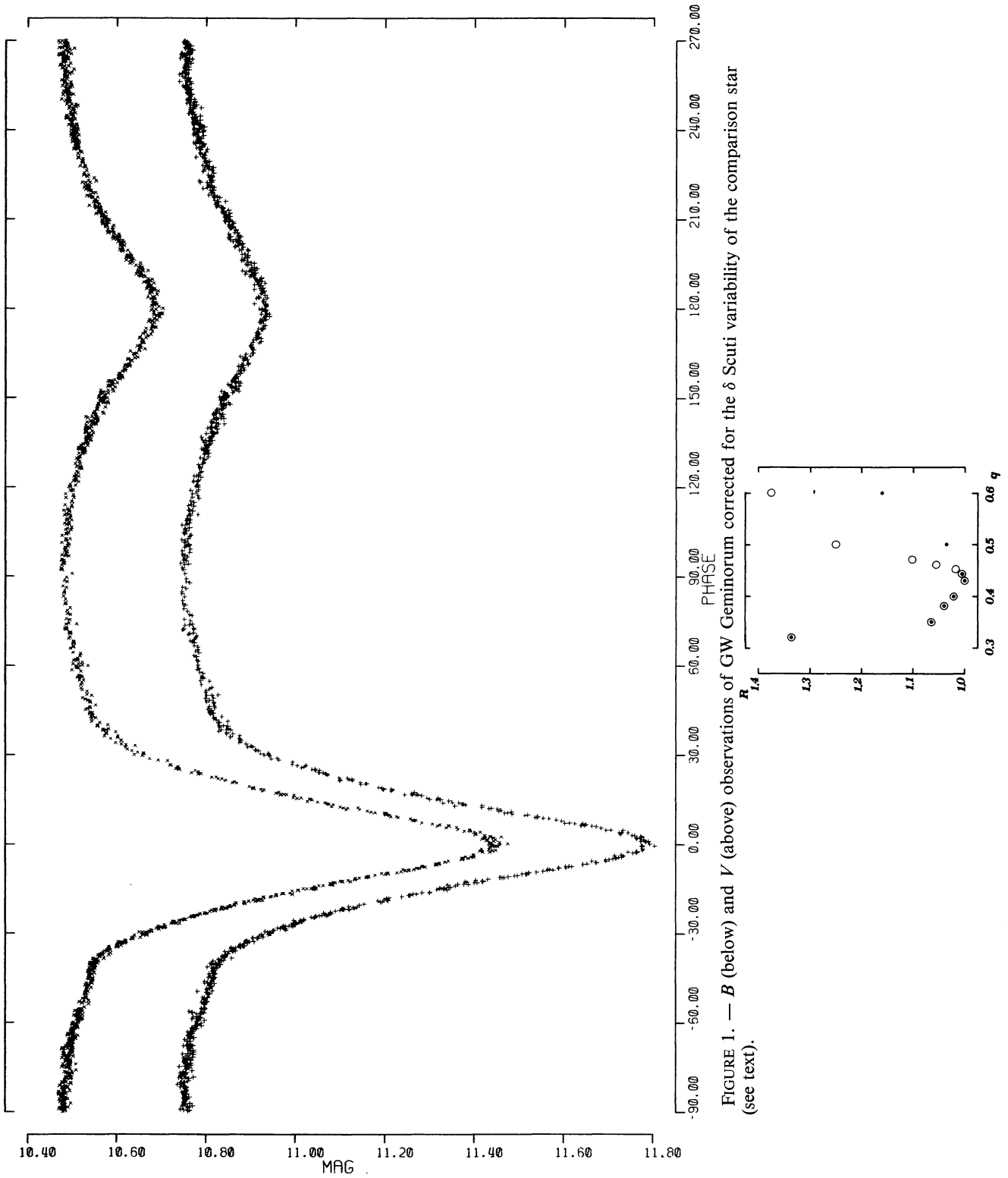


FIGURE 1. — *B* (below) and *V* (above) observations of GW Geminorum corrected for the δ Scuti variability of the comparison star (see text).

FIGURE 2. — Representation of the solutions calculated with assumed values for the mass ratio q . R is the ratio between the sum of the squares of residuals of a given solution and the corresponding value for the best solution.