

## BEAT PHENOMENA IN THE DWARF CEPHEID AE UMa

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AE UMa is an ultrashort pulsating variable. 1900 photoelectric *B* and *V* observations were carried out at the Merate Observatory during 1974. They enable us to determine nineteen new epochs of maximum and nineteen epochs of minimum light and to indicate a strong modulation in the light curves. The main period has been improved:  $P_0 = 0^d 086016883$ . The residuals computed by means of  $P_0$  and the variation in the brightness of the maxima and of the minima can be accounted for by a beat phenomenon whose period is found to be:  $P_b = 0^d 293616$ . The secondary period follows:  $P_1 = 0^d 066527$  and the ratio  $P_1/P_0 = 0.7734$ . This ratio and the amplitudes of the oscillations relating to  $P_0$  and to  $P_1$  are compared with those of other double mode AI Vel variables.

*Key words:* dwarf cepheids – AI Vel stars – beat phenomena

## 1. INTRODUCTION

The ultrashort-period cepheids, or AI Vel stars, or dwarf cepheids studied up to the present are a small group of stars, some of which show remarkable beat phenomena in the light curves. Photometric observations of some variables of this type were started at the Merate Observatory during 1952 and are being continued to check possible secular changes of the pulsation period. Beat phenomena were also shown to be present in the variables RV Ari (Broglia and Pestarino 1955, Broglia 1958) and BP Peg (Broglia 1959).

During the past year, Tsesevich (1973) presented evidence that the variable AE UMa has a very short period ( $P = 124$  min.) with no secular variations, and moreover that its light curve shows large amplitude variations. After this announcement, light variations of this pulsating star were followed during seven nights on February and March 1974. A beat phenomenon appeared at once, and a provisional value for the beat period was evaluated (Broglia and Conconi 1974).

During the first half of 1974 Szeidl (1974), from more than one thousand photoelectric measures obtained at Konkoly Observatory, also determined the beat period and the amplitude of the beat modulation at maximum light. The literature of the variable has also been reviewed by Szeidl (1974).

## 2. THE PHOTOMETRIC OBSERVATIONS

*B, V* measurements, in a system close to *UBV*, were obtained with the single channel photoelectric photometer at the 102 cm Zeiss reflector of Merate Observatory. The output from a charge-integrator system was recorded on a Speedomax chart recorder. The short integration time, twenty seconds, and the rapid switch from variable to comparison star by means of an automatic device, allowed us to reduce the influence of variable atmospheric transmission and to monitor accurately the rapid brightness variations near the maximum. Observations, once digitized on a measuring machine, were corrected for atmospheric extinction and plotted using a Calcomp plotter. Only the corrections for the colour difference between comparison and variable star were significant. The colour of the variable can change two tenths of a magnitude during a cycle, and varies also from one cycle to another through the beat phenomenon. The reduction programme therefore computes a

preliminary colour curve for the variable, then corrects for the colour extinction. The corrections were at most  $0^m007$  for the  $B$  measurements and  $0^m002$  for  $V$ , but were mostly negligible.

The star  $c = \text{BD} + 46^\circ 1882$  was used as a comparison; it was also used by Szeidl. Check stars were  $a = \text{BD} + 46^\circ 1884$  and  $b$ , an eleven magnitude star, south-west of the variable ( $\Delta\alpha = 0.63$  min.;  $\Delta\delta = 2.5$ ). These stars give no sign of any variability. The mean of the magnitude differences between comparison and check stars, and the mean errors for a single  $\Delta m$  were:

$B$	$V$
$m_c - m_a = -0^m386 (\pm 0^m006)$	$-0^m099 (\pm 0^m004)$
$m_c - m_b = -1.423 ( .009)$	$-1.197 ( .006)$

The reliability of our differential photometry of AE UMa can be seen from the above mean errors. A transfer to the international system was achieved by comparing  $c$  with the star no. 1 ( $V = 8^m47$ ,  $B - V = +0^m13$ ) from the field 920-35 (Sanders 1966). We obtained for the comparison star  $c$ :

$$V = 10^m19 \quad B - V = +0^m30$$

The individual observations of AE UMa are listed as a function of heliocentric Julian Date in tables 1 and 2, and they are plotted in figures 1 and 2.

### 3. FUNDAMENTAL AND BEAT PULSATIONS

The light curves have the typical asymmetrical shape of the dwarf cepheids; moreover the large variation in the amplitude of the light curves suggests that AE UMa is a multiperiodic variable. The beat phenomenon is much more evident near maximum light than near minimum. The amplitudes of the colour curves appear to be roughly proportional to the amplitudes of the light curves, and vary in phase.

To study the characteristics of the beat modulation, we have fitted the observations near the top and bottom of the light curves with third order polynomials. A least squares calculation gave us the instants and the magnitudes at maximum and at minimum. The standard error of a single magnitude with respect to the polynomial fitting gave the result  $\pm 0^m005$  for both the colours, comparable to the errors quoted above for the stars  $c$ ,  $b$ ,  $a$ . The values for the maximum are listed in table 3, together with those reported in the literature; the epochs and the magnitudes obtained for the minimum are given in table 4. All the epochs are the mean of the  $B$  and  $V$  results. Moreover, whilst the difference between the  $B$  and  $V$  instants of maximum is never greater than  $0^d0004$  and the differences are randomly distributed, we note a systematic difference, the value of which is on the average  $+0^d0009$ , in the sense minimum  $V$  minus minimum  $B$ . To derive the fundamental period  $P_0$  we considered all the epochs of maximum light. The entire set of observations gave the result:

$$\text{Max.} = \text{Hel.JD } 2442062.5823 + 0.086016883 n$$

The cycle number  $n$  and the residuals  $(O-C)_1$  between observed times of maximum and the times computed from the above expression are given in table 3; the corresponding values for the minimum, with reference to the normal epoch:  $\text{Min.} = \text{Hel.JD } 2442062.5526$ , are listed in table 4. The value for  $P_0$  confirms and slightly improves the one determined by Tsevech (1973). The study of the beat phenomenon relies on photoelectric observations only. The main period  $P_0$  was held constant, and the observations were represented by a sine function:  $A \sin(f_b - f_0) + C$ . Starting from the provisional value for the beat period (Broglia and Conconi 1974),  $P_b$  was varied to minimize the sums of the squares of the residuals. Least squares solutions were obtained for the amplitude  $A$  of the sine curve and for the constants  $f_0$  and  $C$ . The Szeidl  $(O-C)_1$  and our residuals were considered together. The magnitudes at maximum light for both colours were handled in a similar way. However due to a possible small difference between Szeidl's instrumental colour system and ours, the two sets of measurements were treated separately. The same computation was then performed for the magnitudes and the

instants of minimum light. All the values calculated for  $P_b$  agree; the more accurate determination, however, is that obtained from the residuals  $(O-C)_1$ . It gave:

$$P_b = 0.293616 \pm 0.000022 \text{ m.e.}$$

in accordance with the value  $P_b = 0.29364$  given by Szeidl (1974). With this value, sine curves were again fitted to the observations, and beat phases were computed by means of the formula:

$$f_b = (\text{Hel.JD} - 2442062.5825) P_b^{-1}$$

The observations, plotted against beat phases, are shown in figure 3 and figure 4 together with the computed curves. The corresponding residuals,  $(O-C)_2$  for the epochs, and  $(O-C)_B$  and  $(O-C)_V$  for the magnitudes are given in tables 3, 4. The means of the residuals are comparable to the estimated precision of the observations (epochs and magnitudes) noted above. We consider therefore that the observed and computed points agree well, even though a few measurements deviate a little: in our opinion these latter do not justify a possible third period.

#### 4. CONCLUSIONS

Our results call for the following comments, and for a comparison with the AI Vel stars which show beat phenomena.

a) The main period  $P_o$  seems to be constant over the more than one hundred and fifty thousand cycles interval monitored by the observations. In this connection, AE UMa falls in with the dwarf cepheids, which have stable periods unlike the cluster variable (Balazs-Detre and Detre 1966). We believe that the marked residuals  $(O-C)_1$  of some former epochs are incorrect, and are due to relative poor precision, and perhaps to an unfavourable distribution of the old visual and photographic measurements.

b) The skew shape of the light curve, its amplitude and the Ludendorf coefficient for the basic pulsation:  $(\text{epoch max.} - \text{epoch min.})/P_o = 0.35$  are like those found for the AI Vel stars with moderate beat distortion.

c) The main pulsation is characterized by the following parameters, calculated by least squares:

	B	V	B-V
Max.	$11^m 131 \pm 0^m 002 \text{ m.e.}$	$10^m 985 \pm 0^m 001 \text{ m.e.}$	$+ 0^m 146$
Min.	11.774 .002	11.489 .003	+ .285
total amplitude	0.64	0.50	0.14

The colour indices are comprised in the range common to these variables (Bessell 1969).

d) According to the interpretation that the beat phenomenon results from interference of two pulsations, the period  $P_1$  of the first overtone pulsation follows from the relation:  $1/P_1 = 1/P_o + 1/P_b$ . It gives the result:  $P_1 = 0^d 066527$ . In table 5 the periods are given for the AI Vel stars with two excited overtones. The ratio  $P_1/P_o$  for the five variables with  $P_o$  between  $0^d 08$  and  $0^d 15$  are very close, which for the remaining stars, however, is also close to the mean value for the group  $P_1/P_o = 0.773$ .

e) The oscillations of the instants of maximum and minimum light are in phase (figure 3). Their amplitudes gave the same result:  $0^d 0027 \pm .0001 \text{ m.e.}$  and  $0^d 0022 \pm .0004 \text{ m.e.}$  respectively. Amplitudes of the same order were evaluated for RV Ari,  $0^d 0048$  (Broglia 1958) and for BP Peg,  $0^d 003$  (Broglia 1959), whilst for V 703 Sco, which shows a stronger modulation of the light curves,  $A = 0^d 007$  (Ponson 1963).

We note, moreover, that there is a phase lag of about  $P_b/4$  with respect to the sine curve representing the magnitudes at maximum light (figure 4). It appears also (figure 3 and 4) that starting from the phase  $f_b$  when the variable is brightest and the  $O-C$  is near zero, since  $P_1$  is below  $P_o$ , the light maximum resulting from interference of the two pulsations comes earlier and the corresponding  $O-C$  is more negative.

The larger dispersion of  $(O-C)_1$  around the sine curve for the minima is due to the poor precision of the instants computation. The mean deviation, however is only one minute, the same as the precision estimated above.

f) The way that the secondary pulsation modulates the brightness of the maximum and minimum light is shown in figure 4. The following total amplitudes of the modulation were calculated by least squares:

	$B$	$V$	$B-V$
Max.	0 <sup>m</sup> 240	0 <sup>m</sup> 182	0 <sup>m</sup> 06
Min.	.044	.032	.01

The magnitude distortion due to the interfering pulsations appears to be much stronger near maximum than at minimum light. From figure 4, it appears also that the two modulations differ in phase  $P_b/2$ : when the maximum is brightest, the minimum is faintest. An analogous phase-relation holds for SX Phe (Stock and Tapia 1971) and also for AI Vel (Walraven 1955).

g) In the last five columns of table 5, the total amplitudes  $2A$  (for the  $V$  and  $B$  measurements) of the light oscillations with periods  $P_0$  and  $P_1$  respectively, have been listed for all the double mode AI Vel stars. It appears that the modulation can be very different. For AI Vel and V 703 Sco, the two pulsations have about the same amplitude: accordingly the light curves undergo conspicuous variations, and humps also appear. For RV Ari, VZ Cnc, AE UMa and SX Phe the modulation is moderate, it touches mainly the light curves near the maximum and no hump appears. The  $P_1$  pulsation of BP Peg is still smaller; for CY Aqr, the star of the group with the shortest period, the beat phenomenon is very small and moreover little humps can be seen near the minimum. The conclusion follows that whilst the ratios of the observed periods, first overtone to the main pulsation, of all the double mode AI Vel variables are practically the same, the ratios of the pulsational energies of the two excited modes can be very different.

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Table 1 *B* observations of AE UMa

HEL. J.D. 24.....	$\beta$	HEL. J.D. 24.....	B	HEL. J.D. 24.....	B	HEL. J.D. 24.....	B	HEL. J.D. 24.....	B
42065.5363	11.601	42065.6038	11.252	42065.6691	11.433	42068.3567	11.294	42068.4346	11.243
.5389	11.639	.6056	11.286	.6711	11.290	.3612	11.382	.4368	11.275
.5406	11.639	.6070	11.307	.6723	11.195	.3627	11.433	.4388	11.320
.5425	11.662	.6089	11.344	.6731	11.140	.3651	11.476	.4411	11.343
.5433	11.665	.6105	11.372	.6747	11.053	.3658	11.478	.4418	11.347
.5456	11.676	.6122	11.414	.6763	10.991	.3676	11.518	.4436	11.381
.5459	11.690	.6129	11.423	.6769	10.998	.3687	11.526	.4443	11.390
.5506	11.697	.6146	11.455	.6787	10.979	.3708	11.560	.4464	11.420
.5517	11.713	.6154	11.471	.6794	10.999	.3716	11.568	.4473	11.441
.5536	11.712	.6173	11.506	.6820	11.052	.3741	11.603	.4502	11.482
.5543	11.720	.6190	11.524	.6827	11.073	.3756	11.631	.4509	11.488
.5568	11.732	.6197	11.538	.6833	11.103	.3784	11.647	.4528	11.512
.5605	11.734	.6218	11.566	.6852	11.148	.3796	11.656	.4545	11.529
.5620	11.741	.6234	11.581	.6867	11.181	.3822	11.675	.4570	11.553
.5639	11.742	.6250	11.600	.6885	11.230	.3831	11.698	.4581	11.570
.5655	11.749	.6260	11.613	.6892	11.246	.3851	11.717	.4589	11.581
.5673	11.763	.6279	11.628	.6898	11.270	.3912	11.757	.4612	11.617
.5679	11.746	.6286	11.639	.6916	11.315	.3945	11.762	.4620	11.618
.5697	11.749	.6304	11.665	.6933	11.346	.3952	11.778	.4642	11.639
.5711	11.731	.6320	11.690	.6933	11.346	.3952	11.778	.4658	11.647
.5732	11.712	.6339	11.687	.6933	11.346	.3952	11.778	.4678	11.665
.5739	11.714	.6346	11.695	.6933	11.346	.3952	11.778	.4693	11.683
.5743	11.712	.6365	11.708	.6933	11.346	.3952	11.778	.4712	11.694
.5750	11.685	.6372	11.719	.6933	11.346	.3952	11.778	.4721	11.698
.5756	11.694	.6388	11.724	.6933	11.346	.3952	11.778	.4740	11.704
.5786	11.639	.6402	11.733	.6933	11.346	.3952	11.778	.4784	11.726
.5802	11.627	.6420	11.758	.6933	11.346	.3952	11.778	.4803	11.733
.5819	11.576	.6437	11.762	.6933	11.346	.3952	11.778	.4822	11.749
.5826	11.455	.6446	11.765	.6933	11.346	.3952	11.778	.4828	11.753
.5846	11.502	.6469	11.794	.6933	11.346	.3952	11.778	.4851	11.741
.5852	11.457	.6476	11.781	.6933	11.346	.3952	11.778	.4861	11.750
.5869	11.404	.6492	11.792	.6933	11.346	.3952	11.778	.4881	11.756
.5886	11.320	.6507	11.798	.6933	11.346	.3952	11.778	.4899	11.757
.5906	11.235	.6523	11.785	.6933	11.346	.3952	11.778	.4921	11.756
.5913	11.214	.6540	11.792	.6933	11.346	.3952	11.778	.4928	11.763
.5919	11.198	.6558	11.771	.6933	11.346	.3952	11.778	.4949	11.733
.5941	11.178	.6574	11.735	.6933	11.346	.3952	11.778	.4959	11.731
.5948	11.157	.6588	11.724	.6933	11.346	.3952	11.778	.4979	11.721
.5955	11.170	.6613	11.726	.6933	11.346	.3952	11.778	.4997	11.707
.5971	11.171	.6619	11.694	.6933	11.346	.3952	11.778	.5030	11.650
.5978	11.186	.6635	11.682	.6933	11.346	.3952	11.778	.5037	11.640
.5995	11.198	.6652	11.613	.6933	11.346	.3952	11.778	.5057	11.593
.6014	11.214	.6669	11.550	.6933	11.346	.3952	11.778	.5067	11.568
.6032	11.242	.6675	11.523	.6933	11.346	.3952	11.778	.5074	11.548













Table 2  $V$  observations of AE UMa

HEL. J.D. 24.....	$V$	HEL. J.D. 24.....	$V$	HEL. J.D. 24.....	$V$	HEL. J.D. 24.....	$V$	HEL. J.D. 24.....	$V$
42065.5366	11.359	42065.6053	11.104	42065.6694	11.191	42068.3720	11.324	42068.4469	11.226
•5395	11.381	•6061	11.115	•6715	11.082	•3745	11.362	•4476	11.225
•5439	11.378	•6085	11.149	•6726	11.015	•3760	11.374	•4505	11.267
•5479	11.397	•6101	11.167	•6734	10.972	•3789	11.385	•4512	11.269
•5436	11.404	•6118	11.193	•6751	10.923	•3799	11.406	•4533	11.295
•5453	11.407	•6126	11.199	•6766	10.894	•3826	11.414	•4550	11.310
•5462	11.424	•6143	11.231	•6772	10.897	•3835	11.429	•4574	11.326
•5488	11.431	•6150	11.242	•6791	10.904	•3855	11.439	•4584	11.329
•5511	11.431	•6169	11.271	•6798	10.902	•3915	11.485	•4592	11.347
•5533	11.437	•6187	11.274	•6817	10.938	•3948	11.480	•4616	11.364
•5539	11.436	•6193	11.291	•6823	10.940	•3955	11.472	•4623	11.370
•5562	11.454	•6211	11.315	•6830	10.964	•3978	11.488	•4645	11.375
•5590	11.468	•6229	11.330	•6847	10.994	•3991	11.483	•4661	11.395
•5617	11.476	•6247	11.344	•6863	11.024	•4032	11.500	•4690	11.408
•5636	11.475	•6254	11.363	•6881	11.051	•4057	11.487	•4697	11.418
•5651	11.468	•6275	11.382	•6889	11.051	•4064	11.483	•4716	11.420
•5670	11.459	•6283	11.377	•6895	11.079	•4086	11.478	•4725	11.438
•5676	11.472	•6300	11.401	•6913	11.112	•4093	11.460	•4743	11.440
•5693	11.474	•6317	11.415	•6938	11.142	•4112	11.438	•4788	11.449
•5708	11.464	•6335	11.420	•6954	11.140	•4124	11.425	•4806	11.466
•5727	11.457	•6343	11.417	•3311	11.347	•4143	11.390	•4825	11.464
•5736	11.442	•6362	11.433	•3334	11.275	•4162	11.360	•4833	11.465
•5755	11.442	•6369	11.434	•3346	11.236	•4182	11.316	•4856	11.465
•5763	11.436	•6385	11.445	•3359	11.104	•4193	11.283	•4865	11.473
•5782	11.404	•6399	11.447	•3376	11.069	•4200	11.263	•4885	11.482
•5799	11.384	•6417	11.458	•3393	10.992	•4228	11.171	•4903	11.481
•5816	11.336	•6434	11.468	•3400	10.944	•4235	11.155	•4924	11.477
•5823	11.322	•6443	11.463	•3425	10.910	•4241	11.129	•4933	11.480
•5843	11.279	•6462	11.482	•3451	10.916	•4260	11.094	•4953	11.466
•5849	11.266	•6473	11.472	•3457	10.931	•4266	11.077	•4963	11.462
•5866	11.210	•6489	11.468	•3475	10.961	•4278	11.056	•4983	11.445
•5882	11.158	•6503	11.482	•3481	10.966	•4298	11.054	•5001	11.433
•5903	11.076	•6520	11.486	•3499	11.009	•4305	11.046	•5034	11.398
•5909	11.060	•6527	11.469	•3517	11.022	•4313	11.061	•5040	11.382
•5916	11.055	•6554	11.475	•3536	11.060	•4320	11.052	•5059	11.347
•5936	11.019	•6561	11.457	•3544	11.076	•4342	11.072	•5063	11.341
•5945	11.012	•6578	11.453	•3563	11.099	•4349	11.075	•5070	11.322
•5952	10.997	•6592	11.445	•3573	11.122	•4373	11.099	•5089	11.277
•5968	11.017	•6616	11.431	•3616	11.199	•4388	11.126	•5096	11.251
•5974	11.015	•6622	11.443	•3630	11.225	•4396	11.136	•5111	11.211
•5992	11.042	•6641	11.398	•3654	11.245	•4413	11.143	•5117	11.189
•6008	11.047	•6655	11.348	•3651	11.262	•4421	11.159	•5137	11.131
•6027	11.060	•6672	11.297	•3683	11.288	•4440	11.179	•5144	11.105
•6035	11.081	•6678	11.271	•3713	11.314	•4446	11.190	•5151	11.088



Table 2 (continued)

HEL. J.D. 24.....	V	HEL. J.D. 24.....	V	HEL. J.D. 24.....	V	HEL. J.D. 24.....	V	HEL. J.D. 24.....	V
42069.4454	11.453	42069.5141	11.477	42069.5855	11.401	42069.6543	11.253	42086.5259	11.352
4484	11.413	5155	11.490	5879	11.412	6564	11.267	5289	11.386
4507	11.395	5173	11.486	5856	11.430	6581	11.280	5314	11.420
4516	11.364	5196	11.492	5803	11.427	6598	11.290	5350	11.435
4535	11.300	5213	11.493	5919	11.450	6615	11.320	5395	11.463
4555	11.211	5230	11.495	5937	11.460	6631	11.337	5449	11.500
4561	11.196	5237	11.492	5953	11.464	42086.4203	11.130	5499	11.505
4569	11.159	5253	11.486	5988	11.470	4240	11.183	5536	11.517
4588	11.051	5272	11.464	5984	11.471	4277	11.246	5577	11.508
4595	11.023	5293	11.455	6000	11.476	4299	11.287	5617	11.476
4602	10.986	5312	11.421	6019	11.469	4324	11.293	5643	11.398
4618	10.954	5329	11.389	6036	11.476	4358	11.309	5671	11.336
4626	10.931	5346	11.357	6052	11.481	4423	11.392	5685	11.280
4635	10.934	5364	11.290	6073	11.468	4506	11.464	5692	11.256
4657	10.914	5382	11.240	6089	11.474	4557	11.476	5710	11.171
4663	10.921	5397	11.163	6109	11.450	4628	11.504	5718	11.129
4670	10.931	5412	11.094	6124	11.446	4681	11.522	5740	11.014
4690	10.974	5429	11.021	6140	11.445	4725	11.497	5747	10.991
4702	10.992	5436	10.986	6159	11.425	4780	11.471	5762	10.960
4724	11.032	5454	10.953	6176	11.403	4816	11.429	5784	10.951
4732	11.044	5476	10.944	6182	11.400	4852	11.347	5791	10.949
4755	11.103	5460	10.951	6198	11.386	4866	11.232	5810	10.962
4762	11.116	5483	10.953	6215	11.354	4882	11.188	5829	10.984
4777	11.126	5501	10.981	6231	11.337	4895	11.070	5868	11.041
4801	11.175	5525	11.019	6250	11.286	4918	11.070	5900	11.118
4809	11.192	5544	11.045	6268	11.245	4945	10.966	5945	11.191
4816	11.193	5559	11.086	6284	11.202	4945	10.923	5975	11.242
4835	11.227	5568	11.079	6299	11.151	4959	10.916	6008	11.267
4844	11.242	5589	11.117	6316	11.121	4982	10.931	6046	11.295
4868	11.268	5603	11.134	6333	11.097	4988	10.935	42087.3807	11.289
4875	11.274	5622	11.162	6341	11.090	5012	10.961	3823	11.314
4895	11.306	5631	11.186	6357	11.092	5032	11.018	3839	11.332
4903	11.309	5648	11.212	6364	11.076	5053	11.064	3848	11.339
4930	11.347	5667	11.238	6379	11.094	5061	11.076	3865	11.338
4943	11.352	5676	11.239	6388	11.089	5082	11.108	3871	11.352
4963	11.368	5696	11.267	6404	11.101	5100	11.153	3889	11.376
4976	11.383	5714	11.281	6411	11.118	5118	11.173	3903	11.382
4997	11.405	5732	11.311	6426	11.119	5148	11.223	3926	11.397
5018	11.409	5750	11.337	6443	11.134	5155	11.226	3932	11.401
5058	11.427	5765	11.328	6461	11.151	5177	11.275	3953	11.426
5062	11.439	5783	11.368	6476	11.175	5185	11.289	3971	11.435
5080	11.447	5801	11.341	6494	11.196	5204	11.324	3984	11.447
5099	11.472	5818	11.373	6511	11.216	5220	11.349	4002	11.463
5122	11.478	5836	11.401	6528	11.237	5239	11.336	4008	11.456







Table 3 Epochs and magnitudes at the maximum light

Observer	Hel. J.D.	n	(O-C) <sub>1</sub>	(O-C) <sub>2</sub>	B	(O-C) <sub>B</sub>	V	(O-C) <sub>V</sub>
T	28632.398	-156134	- <sup>d</sup> .025					
F	31875.122	118436	+0.035					
F	33379.256	100949	-0.008					
F	35601.188	75118	+0.022					
T	35604.337	75081	-0.012					
F	35607.173	75048	-0.014					
F	35981.202	70700	+0.013					
T	38106.402	45993	-0.006					
T	41059.368	11663	+0.001					
T	41773.223	-3364	+0.001					
S	42062.5835	0	+0.0012	+ <sup>d</sup> .0007	11 <sup>m</sup> .26	+ <sup>m</sup> .02	11 <sup>m</sup> .09	+ <sup>m</sup> .02
B,C	65.5956	+35	+0.0027	-0.0002	11.165	+0.007	.006	+0.002
B,C	.6781	36	-0.0008	-0.0001	10.984	-0.027	10.890	-0.003
F,C	68.3438	67	-0.0016	+0.0002	11.029	+0.005	.903	.000
B,C	.4303	68	-0.0011	-0.0002	.217	+0.004	11.049	+0.002
B,C	.5203	69	+0.0028	-0.0001	.187	+0.002	.024	-0.001
B,C	.6029	70	-0.0006	-0.0005	.016	+0.001	10.899	+0.002
B,C	.6871	71	-0.0024	-0.0001	.130	+0.003	.987	+0.006
B,C	69.3808	79	+0.0032	+0.0006	.214	-0.001	11.050	-0.006
B,C	.4651	80	+0.0014	+0.0005	.036	-0.001	10.917	+0.004
B,C	.5473	81	-0.0024	+0.0001	.071	-0.005	.942	-0.001
B,C	.6363	82	+0.0006	-0.0002	.253	+0.003	11.085	+0.010
B,C	86.4965	278	+0.0015	+0.0006	.046	+0.012	10.913	+0.002
B,C	.5787	279	-0.0023	+0.0002	.081	+0.002	.944	-0.002
B,C	87.4390	289	-0.0022	.0000	.039	+0.001	.904	-0.010
B,C	.5263	290	-0.0009	-0.0006	.222	-0.009	11.049	-0.011
B,C	.6155	291	+0.0023	-0.0006	.154	-0.007	.004	-0.002
S	95.5293	383	+0.0025	-0.0003	.18	-0.02	11.04	-0.01
S	.6118	384	-0.0010	-0.0010	.00	-0.01	10.88	-0.01
B,C	103.3513	474	-0.0030	-0.0012	.182	+0.011	11.024	+0.009
S	106.4520	510	+0.0011	-0.0003	.03	-0.02	10.92	.00
S	119.5258	662	+0.0003	+0.0010	.18	-0.03	11.02	-0.02
S	121.5017	685	-0.0022	+0.0001	.01	-0.02	10.89	.00
B,C	122.3628	695	-0.0012	+0.0003	.028	+0.010	.901	+0.002
B,C	.4484	696	-0.0016	-0.0004	.187	-0.011	11.032	-0.004
S	128.2968	764	-0.0024	-0.0002	.12	.00	11.00	+0.03
S	.3872	765	+0.0020	-0.0002	.25	+0.01	11.10	+0.02
S	.4727	766	+0.0015	.0000	.07	+0.01	10.96	+0.03
S	.5550	767	-0.0022	+0.0001	.06	+0.02	.93	+0.03
S	133.4622	824	+0.0020	+0.0004	.07	+0.01	.92	-0.01
S	.5440	825	-0.0022	+0.0001	.03	.00	.88	-0.01
S	134.4055	835	-0.0009	+0.0008	.02	+0.01	.89	+0.01
S	147.3935	986	-0.0014	+0.0004	.15	.00	.98	-0.01
S	148.4295	998	+0.0024	+0.0004	.09	+0.01	.92	-0.02
S	.5096	999	-0.0036	-0.0014	.03	+0.01	.88	.00
S	159.4365	1126	-0.0008	+0.0009	11.16	.00	.97	-0.03
S	161.4145	+1149	-0.0012	+0.0004	10.99	-0.01	10.86	-0.01

Table 4 Epochs and magnitudes at the minimum light

Hel. J. D.	n	(O-C) <sub>1</sub>	(O-C) <sub>2</sub>	B	(O-C) <sub>B</sub>	V	(O-C) <sub>V</sub>
42065.5652	35	+ <sup>d</sup> .0020	+ <sup>d</sup> .0008	11 <sup>m</sup> .751	- <sup>m</sup> .007	11 <sup>m</sup> .472	- <sup>m</sup> .005
.6502	36	+ .0010	- .0005	.791	- .001	.478	- .024
68.3995	68	- .0023	- .0001	.777	+ .005	.495	+ .008
.4896	69	+ .0018	+ .0012	.757	+ .002	.479	+ .004
.5764	70	+ .0026	+ .0008	.780	- .009	.492	- .008
.6592	71	- .0006	+ .0009	.777	- .010	.489	- .009
69.3480	79	+ .0001	+ .0004	.761	+ .009	.469	- .004
.4360	80	+ .0020	- .0001	.774	- .006	.493	- .001
.5203	81	+ .0003	+ .0011	.790	- .003	.496	- .007
.6021	82	- .0039	- .0022	.758	- .002	.478	- .001
86.4655	278	+ .0002	- .0019	.797	+ .017	.521	+ .027
.5509	279	- .0004	+ .0004	.810	+ .017	.525	+ .022
87.4116	289	+ .0001	+ .0001	.796	.000	.510	+ .005
.4950	290	- .0025	- .0004	.767	.000	.491	+ .007
.5845	291	+ .0010	- .0002	.755	- .002	.482	+ .005
122.4185	696	- .0019	+ .0003	.770	- .005	.481	- .009
.5059	697	- .0005	- .0008	.751	- .002	.464	- .010

Table 5 AI Vel stars with two excites modes

Ref*	Star	P <sub>0</sub>	P <sub>b</sub>	P <sub>1</sub>	P <sub>1</sub> /P <sub>0</sub>	2A <sub>V</sub>	2A <sub>B</sub>	A <sub>0</sub> /A <sub>1</sub>
						(unit=0 <sup>m</sup> .01)		
1	SX Phe	0 <sup>d</sup> .05496437	0 <sup>d</sup> .192834	0 <sup>d</sup> .042773	0.7782	55	21	2.6
2	CY Aqr	.0610383	.17766	.04543	.7443	72	12	6.0
	AE UMa	.08601688	.293616	.066527	.7734	50	21 64	28 2.3
3	RV Ari	.09313	.31634	.07195	.7726	49.5	31 63.5	40 1.6
4	BP Peg	.10954	.3698	.0845	.771	43	9 57	14 4.4
5	AI Vel	.111574	.379188	.086208	.7727	39	51	.76
6	V 703 Sco	.14996	.497295	.115216	.7683	40	50	.80
7	VZ Cnc	.17836376	.716292	.142800	.8006	50	29 70	38 1.8
8	VX Hya	.223389	.761475	.17272	.7732			

\*References: 1. Stock, Tapia (1971); 2. Elst (1972); Fitch (1973); 3. Broglia (1958); 4. Broglia (1959); 5. Walraven (1955); 6. Ponsen (1963); 7. Spinrad (1960); 8. Fitch (1966).

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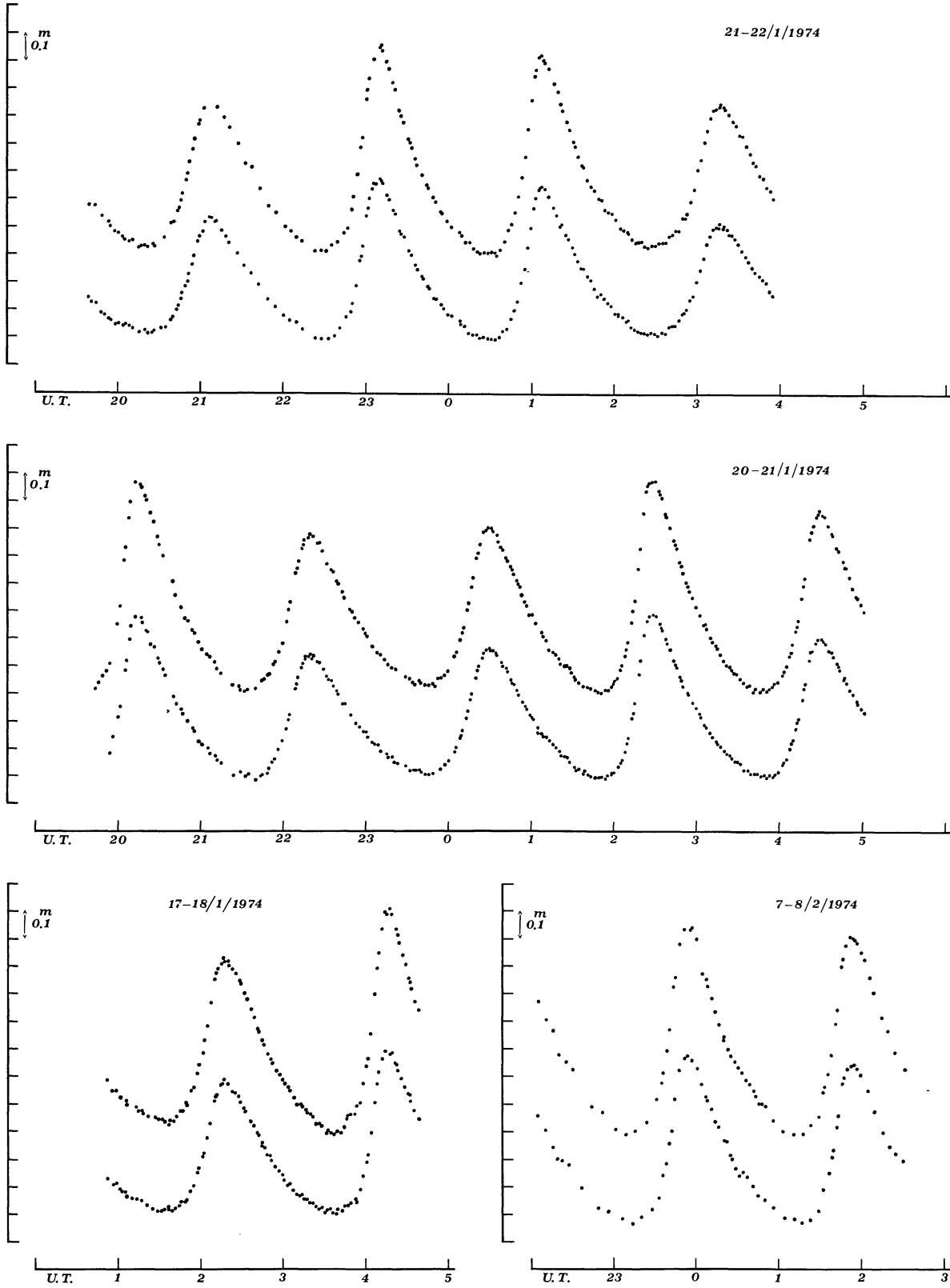


Figure 1 *B* and *V* light curves of AE UMa. The two curves are shifted by the same arbitrary quantity along the vertical axis.

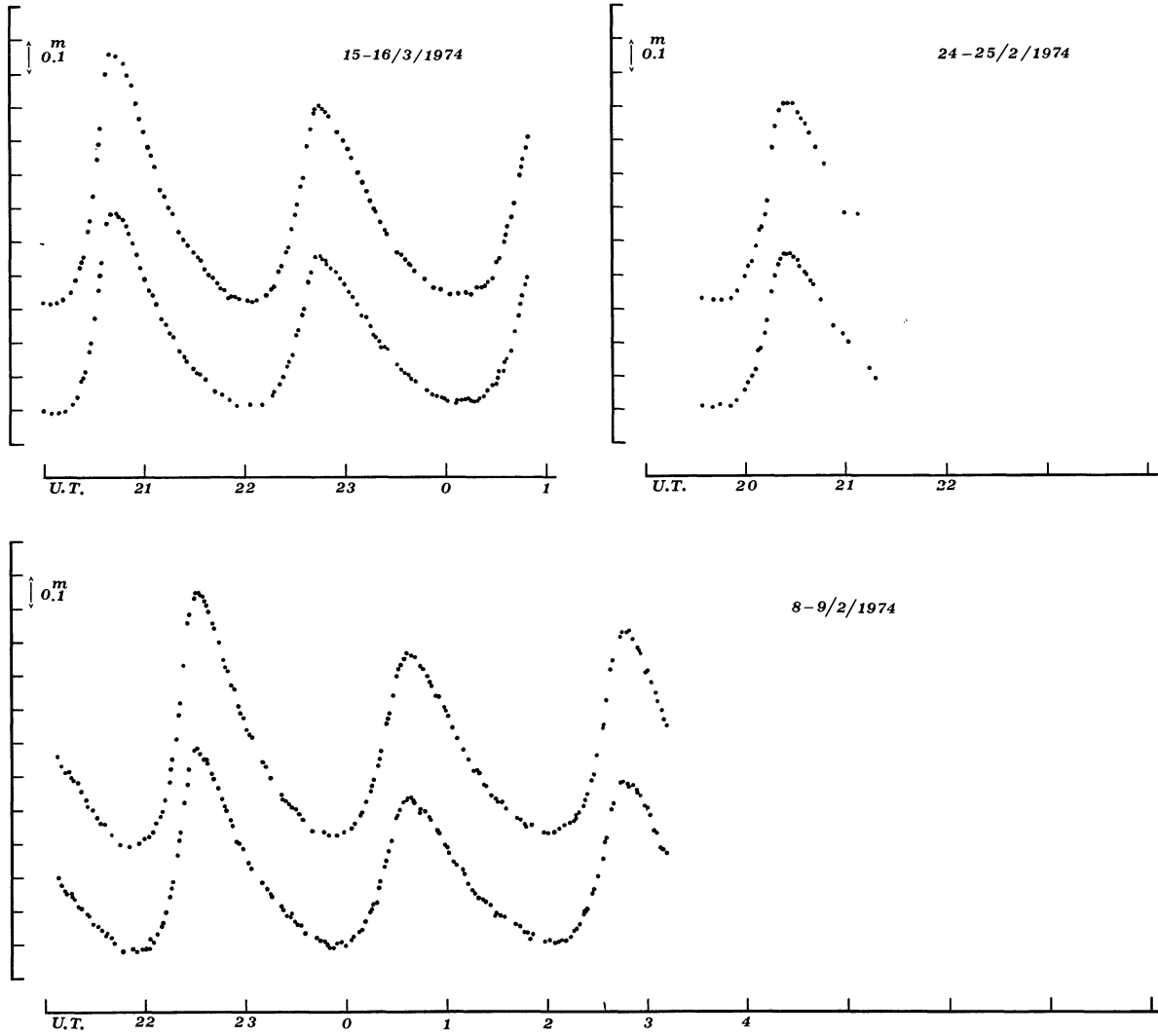


Figure 2 *B* and *V* light curves of AE UMa. The two curves are shifted by the same arbitrary quantity along the vertical axis.



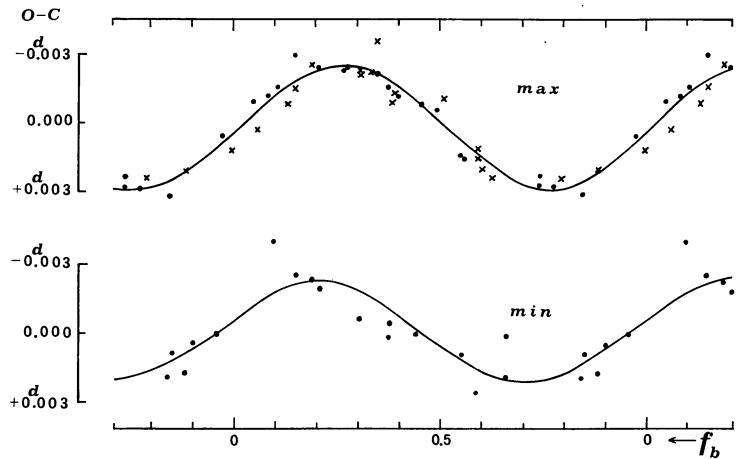


Figure 3 The  $(O-C)_1$ , computed by means of the fundamental period  $P_0$  for all the photoelectric measures, plotted against beat phases. The crosses represent the Szeidl observations, the dots ours. The sinusoids have been calculated by least squares.

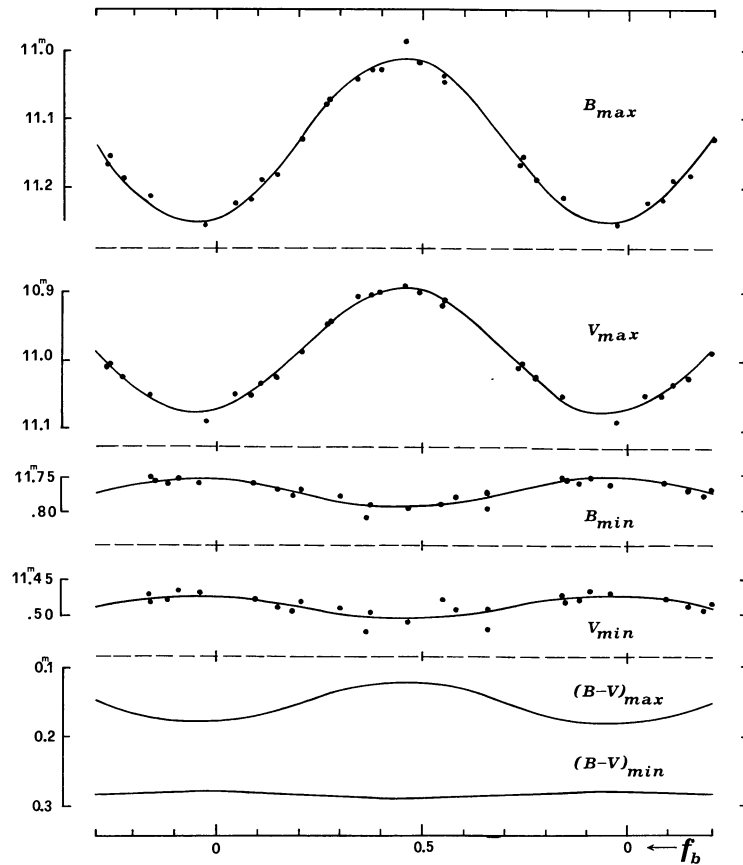


Figure 4 The  $B$ ,  $V$  magnitudes and the colours of AE UMA, at maximum and minimum light, plotted against beat period phases. A  $P_b/2$  phase difference appears between the minimum and maximum curves. The curves have been calculated by least squares.